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Sutton Double-Star Observations.—Under this title the veteran observer, Dr. W. Doberck, publishes in *Astronomische Nachrichten,* Nos. 4394-95, a continuation of his long series of double-star measures. Dr. Doberck is also continuing his investigations of orbits, as appears from two papers published in *Astronomische Nachrichten,* No. 4400, entitled "Synopsis of Elements of Orbits of Some Double Stars," and "On the Orbit of 40 o^2 *Eridani.*" It is to be hoped that Dr. Doberck will soon publish in collected form the results of his extensive investigations in this field.

Among other recent papers of interest to double-star observers we note the publication in *Astronomische Nachrichten,* No. 4406, of a list of eighty-two new double stars by M. Robert Jonckheere, of the observatory at Hem, France. The stars for the most part, are fainter than 9.0 B. D. magnitude, though several of the later discoveries are brighter than this. It is a satisfaction to note that M. Jonckheere has apparently set himself rigorous limits in defining a double star. With one exception, the measured distances are less than 4″, and more than half are less than 2″.

Attention should also be called to a very interesting theoretical paper by Mr. H. C. Plummer, entitled "On the Analytical Method of Finding the Orbit of a Double Star," which is published in the *Observatory* for May, 1910. In this paper Mr. Plummer derives relations between the coefficients of the equation of the apparent ellipse and the geometrical elements of the true orbit, which are, in effect, equivalent to those employed in the well-known Kowalsky method, though by no means identical. Mr. Plummer notes that his method has "points of contact with known graphical methods," and, incidentally, directs attention to a convenient and accurate graphical method published by him in Volume LX (page 483) of the *Monthly Notices, R. A. S.* R. G. Aitken.

June 14, 1910.

NOTE ON THE DOUBLE STAR OΣ 251.

The double star OΣ 251 was examined with the 36-inch telescope on the night of June 2, 1909, and found to be a close, unequal pair. No further measures were obtained until recently, when the earlier observation was confirmed on two nights. March 5 and May 27, 1910. The three measures give:—

1909.420	5°.3	0″.21	1 night
1910.288	14 .0	0 .22	2 nights

As the two components differ a full magnitude in brightness, there is no question as to the correctness of the quadrant.

HUSSEY, in his catalogue of the Otto Struve stars, and BURNHAM, in his General Catalogue, pronounced this star probably single, as it had not been seen double in any large telescope except on one night in 1875, when HALL and HOLDEN, using the Washington 26-inch, suspected elongation in 70° or 80°. BURNHAM failed to see it double with the 18½-inch telescope in 1879, as did SCHIAPARELLI, in Milan, in 1881, 1888, 1889, and 1890. In 1898 and 1899 it was examined by HUSSEY and myself with the 36-inch under excellent conditions and showed no evidence of duplicity. The early measures by OTTO STRUVE, from 1843 to 1853, gave discordant angles ranging from 105° to 152°.

In view of these observations, positive and negative, there seems to be no doubt of the fact that OΣ.251 is a binary star in comparatively rapid orbital motion. It would certainly have been seen double in 1899 if the components then had been as well separated as they are now.[1] R. G. AITKEN.

June 9, 1910.

THE AQUARID METEORS.

On the morning of May 4th, while working at the Crossley dome, quite a number of meteors were seen in the southeast by both Dr. CURTIS and myself. Remembering that the η Aquarid meteors, to which most of these obviously belonged, came to a maximum about this date and were supposed to be connected with Halley's Comet, I began to plot the paths of all seen. As a consequence, a good radiant was secured from six meteors, all quite well plotted. Observations on May 5th failed to supply more than two meteors that were certainly members of the stream. On May 11th, however, another

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[1] OΣ 251 is B_0 D. + 32°.2250, 7.9 magnitude. B_0 D. + 32°.2252, 7.4 magnitude, is 56s following and 2′ north. It is, of course, possible that the latter star was examined on some occasions instead of the right star.